

GAIA DR3 2059297390930630400:

AN ECCENTRIC BINARY WITH CONSIDERABLE APSIDAL MOTION

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Abstract: Apsidal motion in binary stars refers to the gradual precession of the elliptical orbits of two stars around their common center of mass. This phenomenon occurs due to gravitational interactions, which cause the orientation of the orbit to shift over time. We present the indication of apsidal rotation in the eclipsing binary system Gaia DR3 2059297390930630400 (ZTF J195804.88+343708.1) by using the photometric data from the Zwicky Transient Facility (ZTF) and Transiting Exoplanet Survey Satellite (TESS). We used Perando to derive the orbital period from ZTF and times of minima from the TESS data.

1 Introduction

Gaia DR3 2059297390930630400 is an eccentric binary system positioned at RA 19 58 04.88 DEC +34 37 08.026 (J2000). The object has an effective temperature of 7514.8 K according to Gaia DR3 and can therefore be classified as spectral type A/F, which applies to both components of the binary system, as the minima are of comparable depth. In this paper, we present the evidences to show the indication of apsidal motion in this system by obtaining the Times of Minima (ToM). ToM refers to the specific moment/time when the brightness of an eclipsing binary star system is at its lowest. This search is based on Chen et al. (2020), in which we identified the object of interest during a systematic search for eclipsing binaries with asymmetrically located secondary minima.

Similar work with a greater understanding of three southern close eccentric eclipsing binaries has been conducted by Wolf et al. (2024) on GM Nor, V397 Pup, and PT Vel.

2 Data and analysis

We collected photometric data from ZTF (Bellm et al., 2019), obtaining magnitudes in the zg, zi, and zr filters, as well as data from sectors 14, 41, 54, 55, 74, and 75 from the TESS satellite (Ricker et al., 2015). TESS (Transiting Exoplanet Survey Satellite) is mainly dedicated to searching for exoplanets using the transit method, while Zwicky Transient Facility (ZTF) is a ground-based observatory that monitors variable celestial objects. Our analysis was primarily done with Perando - Light Curve and Period Analysis

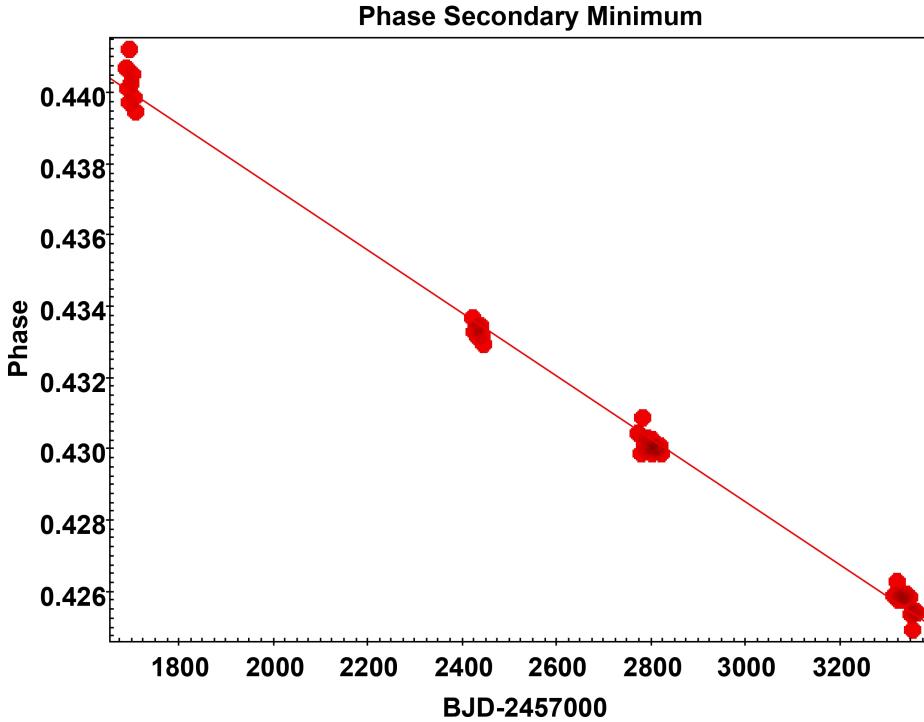


Figure 1: Phase as a function of Secondary Times of Minima

Software (Paunzen & Vanmunster, 2016), including period analysis and analysis of Times of Minima. First, we obtained the system's orbital period as 3.137204 days from the ZTF data, which is in line with the period found in Chen et al. (2020). To refine the period and calculate the ephemeris, we extracted the Times of Minima from each sector of the TESS data, presented in Table 1. Furthermore, the Times of Minima can serve to determine the apsidal motion rate of an eclipsing binary system (Giménez & Bastero, 1995).

Using the method of least squares, the following ephemeris for the primary minima were derived from the Times of Minima:

$$BJD_{minI} = (2458684.7210 \pm 0.0003) + (3.1372013 \pm 8) * E \quad (1)$$

where E is the epoch/cycle number.

Based on this ephemeris, we calculated the phase difference of each of the minimas of all TESS sectors using,

$$\phi = \frac{T - T_0}{P_{orb}} \quad (2)$$

where P_{orb} is the orbital period of the system.

Our primary analysis included plotting all phase values against the Minima Times of the Secondary eclipse as seen in Figure 1.

We have also made a O-C plot using obtained times of minima shown in Figure 2. The plot was built using the ephemeris above.

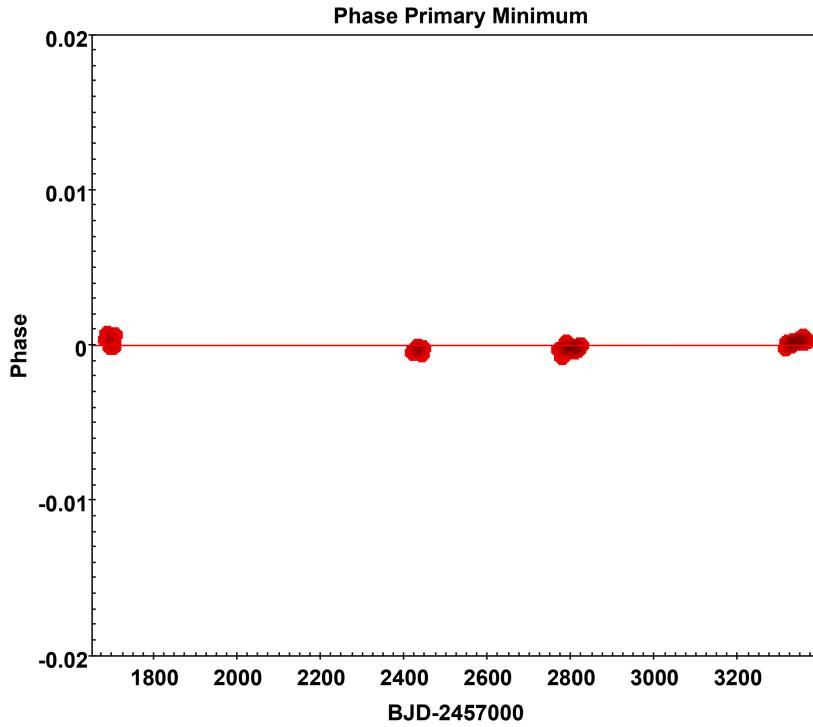


Figure 2: Ephemeris of all primary minima

Figure 3 shows the folded light curve of TESS Sector 14 and Sector 75, showing the variation between the first and the last sector of observations, adding to the evidence of apsidal motion.

3 Conclusion

In conclusion, we observe that a trend is being followed concerning the phase of the secondary eclipse against the secondary minima times. This shows an indication of apsidal rotation in the system. However, we have only observed a small part of the apsidal cycle in the period between BJD 2458684 and BJD 2460366 (1682 days). This aligns with the long periods (about 80, 335, and 160 years) found for the three objects studied in Wolf et al. (2024). Additional observations are required for a better and more complete understanding of the system's behavior.

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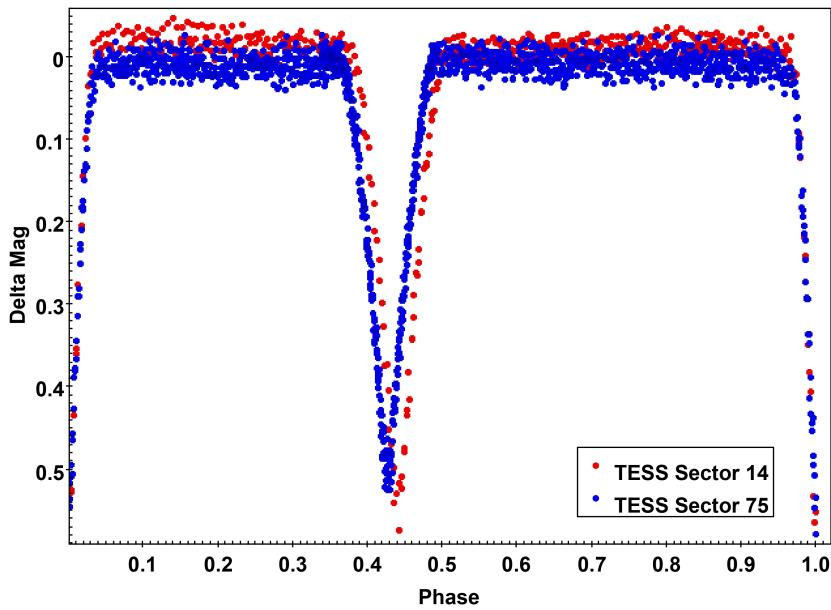


Figure 3: Folded light curve of TESS Sector 14 and Sector 75

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